Figures 2 and 3 show that our kinematical measurements reach the record distance, for absorption spectra, of $\sim 2\ r_e$; this is twice the size of the effective isophote (the one which encircles half of the total light of the galaxy), and corresponds of a surface brightness $\mu_B \simeq 24$.

In order to appreciate the improvement of these over the literature data on absorption line kinematics, one can see for example the "Atlas of velocity dispersion profiles and rotation curves for elliptical and lenticular galaxies" by Di Martino, Busarello and Longo (1989, preprint).

In all cases, our rotation curves exhibit the standard signatures of S0 galaxies, i.e. the bumps characteristic of systems where the "conspiracy" between the disk and the bulge is not finely tuned (cf. Capaccioli, 1979). More importantly, they confirm that the flat trend noted at small galactocentric distances – where disks may be still prominent – is maintained at larger distances, well in the range where the bulge is dominant; for instance, according to Capaccioli et al. (1987), the major axis light profile of NGC 3115 is dominated by the bulge at all galactocentric distances $>0.45\ r_e$. In spite of the fact that, thanks to the unique performances of EFOSC, our measurements have almost doubled the range of radial distances covered by previous kinematical mapping, these deep data do not provide us with any strong evidence for dark matter in early-type disk galaxies, say comparable to the HI rotation curves in late-type disk galaxies. The reasons are the still "short" radial range, and the uncertainties introduced by the non-negligible velocity dispersion and by the modelling. All this matter is under investigation. A full account of the observations described here, and the relative astrophysical discussion, will be presented elsewhere.

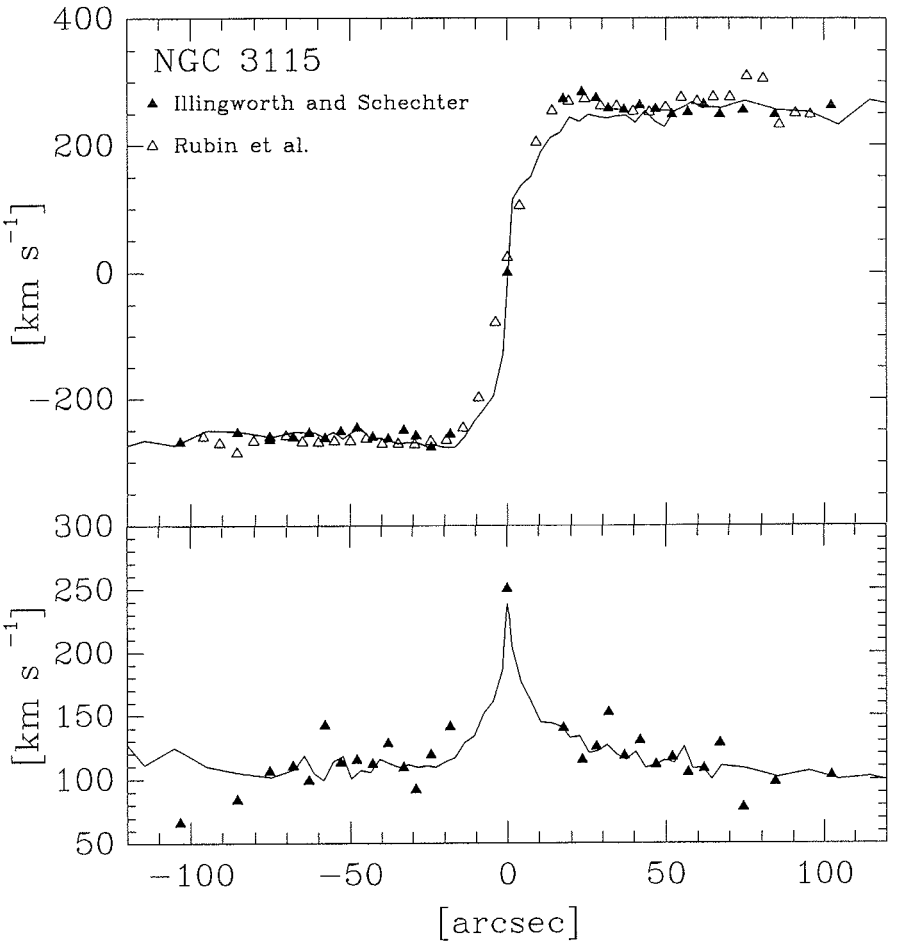

Figure 4: *Comparisons between the rotation and velocity dispersion curves of NGC 3115 from this study (solid line) and the kinematical data published by Rubin et al. (1980) and Illingworth and Schechter (1982).*

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International Portrait Catalogue

One of the last days in September this year, I went to Berlin (West) to give a talk at the Wilhelm Foerster Volkssternwarte (People's observatory), on the occasion of the 100th anniversary of its "Bamberg" refractor.

Among the speakers was also Dieter B. Herrmann, Director of the Archenhold-Sternwarte in Berlin-Treptow, GDR, and noted astronomical historian, who had come over to celebrate this jubilee of one of the oldest, still functioning large telescopes, in this case exclusively installed for public education purposes.

Quite apart from the happy event which brought us together, I learned about Professor Herrmann's efforts to establish the world's most comprehensive collection of portraits of astronomers, in particular for the benefit of (future) historians. More than 7500 photos and drawings have already been gathered by this project which has been going on since 1971, under the auspices of IAU commission 41 (History of Astronomy).

However, it appears that many astronomers are unaware of the existence of this catalogue, that is at least my impression after having talked to some colleagues here in Munich. This is confirmed by Prof. Herrmann's difficulties in acquiring portraits of now living astronomers. He therefore asked whether it would be possible to place a small note in the ESO *Messenger* to this effect, at the same time hoping for reaction from our readers. I am happy to provide this space for this useful purpose.

So Messenger readers beware: here is your best chance ever to make your image known to future generations! I am told that portraits can be sent as gifts or on loan only; in the latter case, they will be returned within a month after having been copied. Subject to any restrictions imposed by the donors, the Archenhold Observatory is prepared to supply on request copies of portraits held by them and will issue lists of their holdings from time to time. That is a most useful offer for all who have to write an article or give a talk about the past of their observatory, etc.

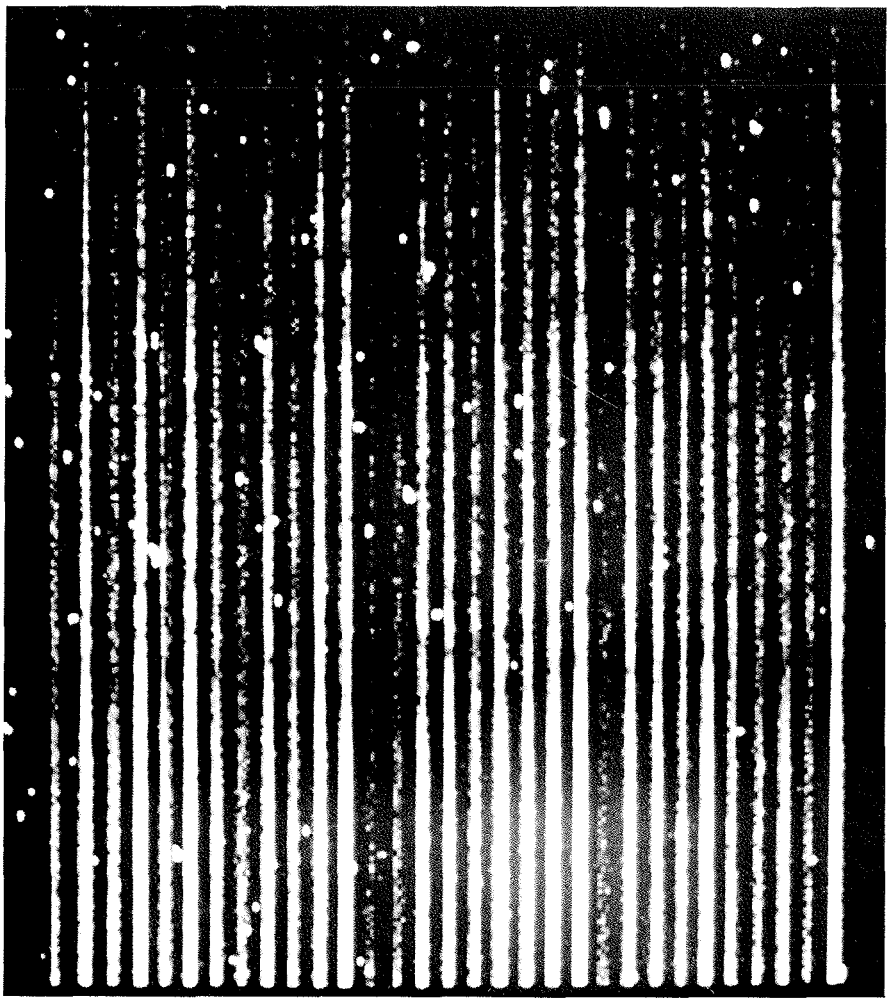
The address is: Prof. Dr. D.B. Herrmann, Archenhold-Sternwarte, Alt-Treptow 1, DDR-1193 Berlin, German Democratic Republic. They will be happy to hear from you. *The editor*

Jacques Beckers Elected to Dutch and Norwegian Academies

On August 28, Queen Beatrix of the Netherlands confirmed the election of ESO staff member Jacques Beckers as "Correspondent" of the Division for Sciences of the Royal Academy of Sciences of the Netherlands. Corresponding members are researchers with a degree from a university in the Netherlands, residing abroad.

Jacques Beckers, who joined ESO in 1988 to become Head of the Interferometry Group, also became a foreign member of the Mathematics-Physics Sciences Division of the Norwegian Academy of Sciences last year.

Our best congratulations to Jacques at the time of these well-deserved honours!



The Efficiency of OPTOPUS

This is part of an OPTOPUS frame resulting from an exposure of galaxies in the cluster Abell 3158, as obtained during the first run (in September 1989) of the ESO Key Programme on "Structure and Dynamics of Rich Clusters of Galaxies".

Note that only the blue parts of the 31 spectra (of 28 galaxies and 3 "skies") are shown in this picture, which covers the wavelength range from ~383 nm (top) to ~440 nm (bottom). The bluest CaII doublet (about one third of the way down) is from sky; the redder CaII doublet (about two thirds of the way down) is from the galaxies, and visually displays the dispersion of the radial velocities of the galaxies in the cluster.

In total, 37 exposures were obtained during the run, which yielded a total of about 1000 galaxy spectra. *P. Katgert (Leiden)*

An Accurate Wavelength Calibration of CCD CASPEC Echelle Spectra

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Introduction

In December 1986 and January 1988 we obtained with CASPEC at the 3.6-m telescope spectra of early-type member stars of the young stellar cluster NGC 2244 with the main purpose of determining accurate radial velocities. We focused our attention on the blue wavelength region (3700–4700 Å) and used CCD # 3 in combination with the 52 l/mm echelle grating.

For the study of the internal kinematics of the cluster and dealing with relatively few lines in early-type stars, it is evidently of primary interest to achieve high precision in the wavelength calibration, particularly avoiding systematic errors with wavelength. Th-Ar calibra-