

Classical Cepheids as tracers of young Galactic and Magellanic stellar populations

R. da Silva ⁽¹⁾, G. Bono ^(1,2), B. Lemasle ⁽³⁾, F. Primas ⁽¹⁾, L. Inno ^(1,5),
K. Genovali ⁽¹⁾, N. Matsunaga ⁽⁴⁾, M. Romaniello ⁽⁵⁾, P. François ⁽⁶⁾,
M. Groenewegen ⁽⁷⁾, and the DIONIZOS collaboration

⁽¹⁾ Università di Roma Tor Vergata, ⁽²⁾ Osservatorio Astronomico di Roma,
⁽³⁾ Astronomical Institute Anton Pannekoek, ⁽⁴⁾ Tokyo University, ⁽⁵⁾ ESO,
⁽⁶⁾ Observatoire de Paris, ⁽⁷⁾ Royal Observatory of Belgium



UNIVERSITA' degli STUDI di ROMA
TOR VERGATA

Introduction

- ★ Classical Cepheids are fundamental objects to estimate Galactic and extragalactic distances

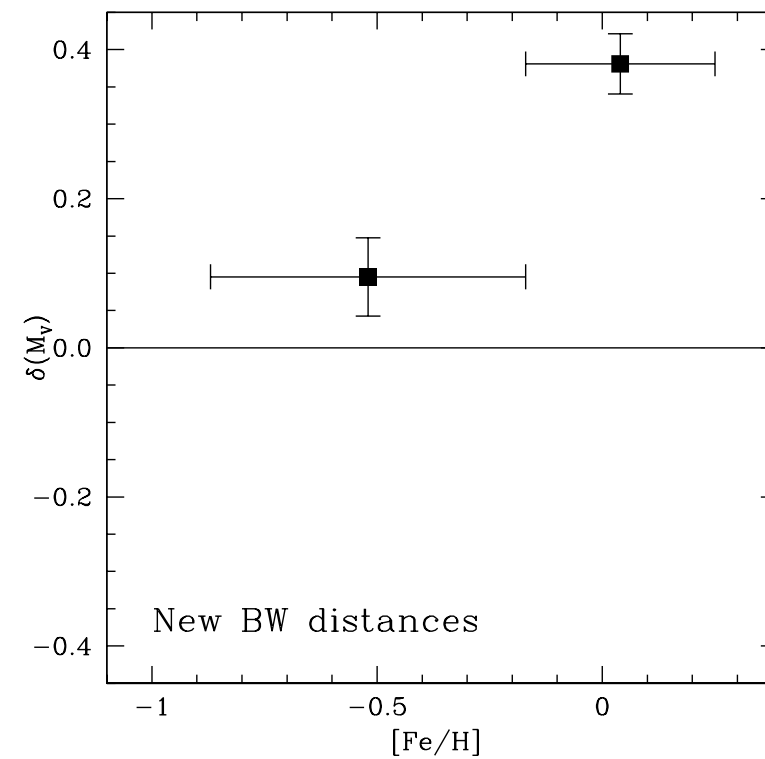
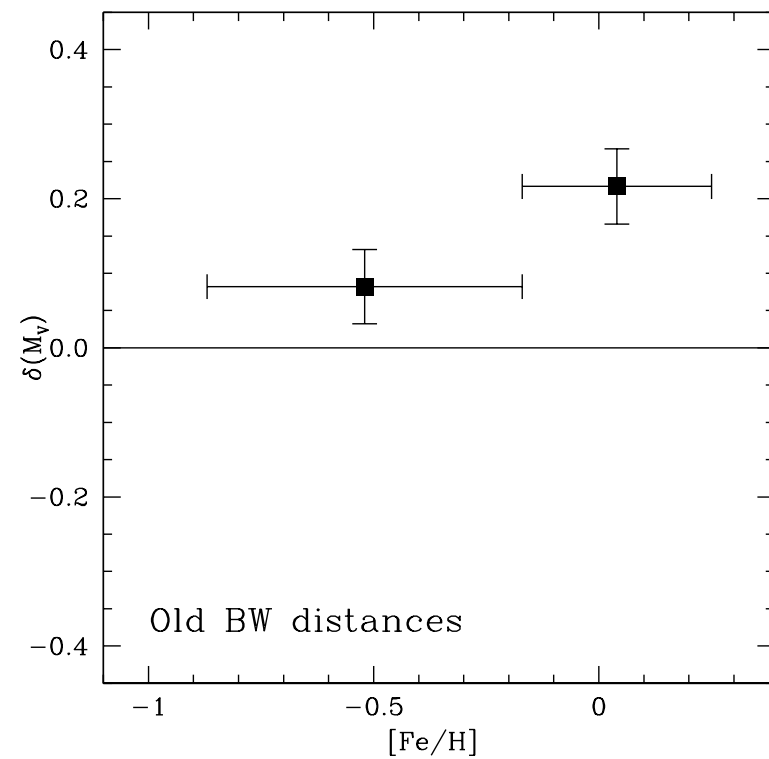
Period-Luminosity relation

What is the influence of the chemical composition on the pulsation properties?

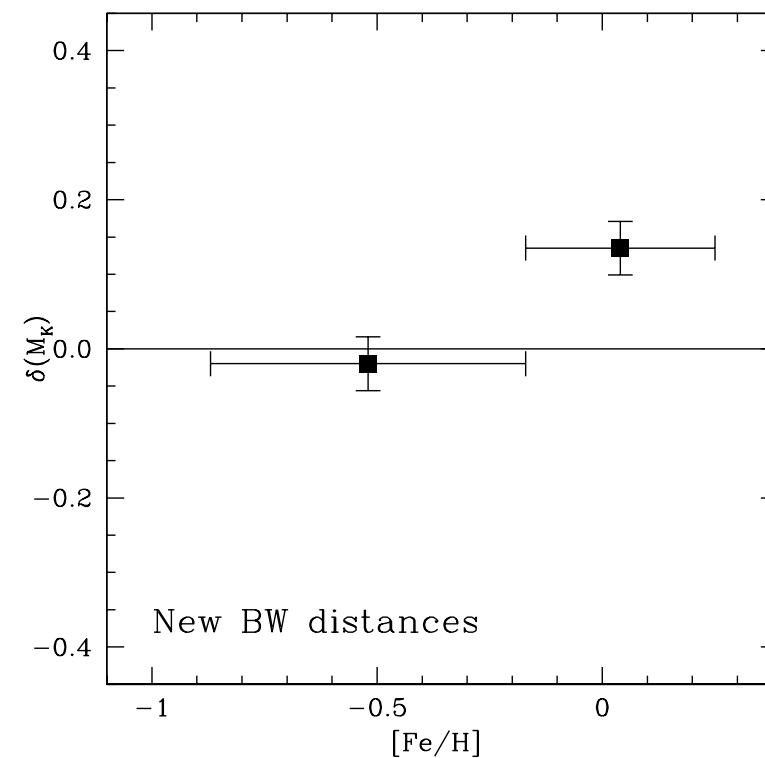
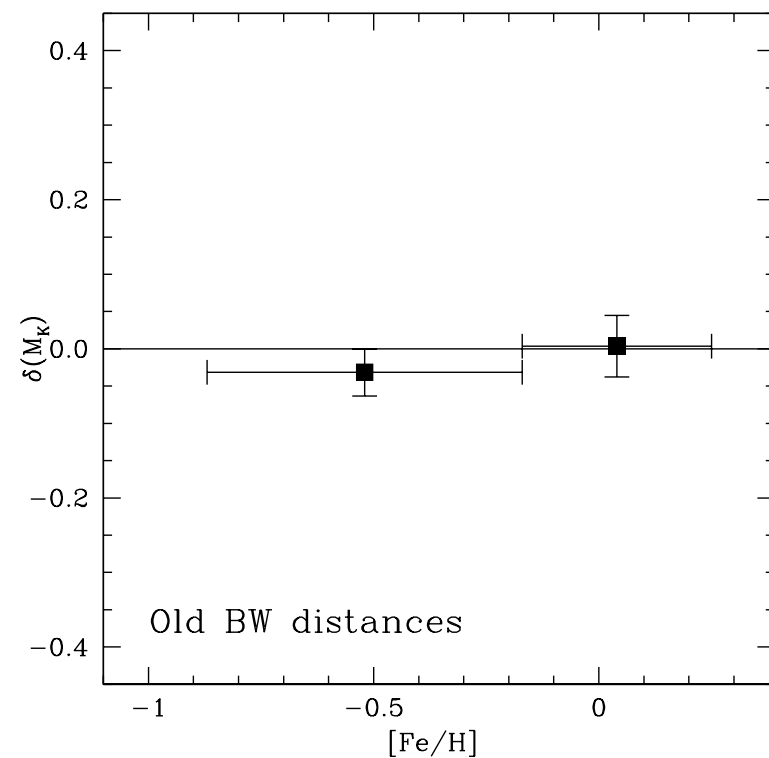
Introduction

Dependence
on metallicity

δM_V



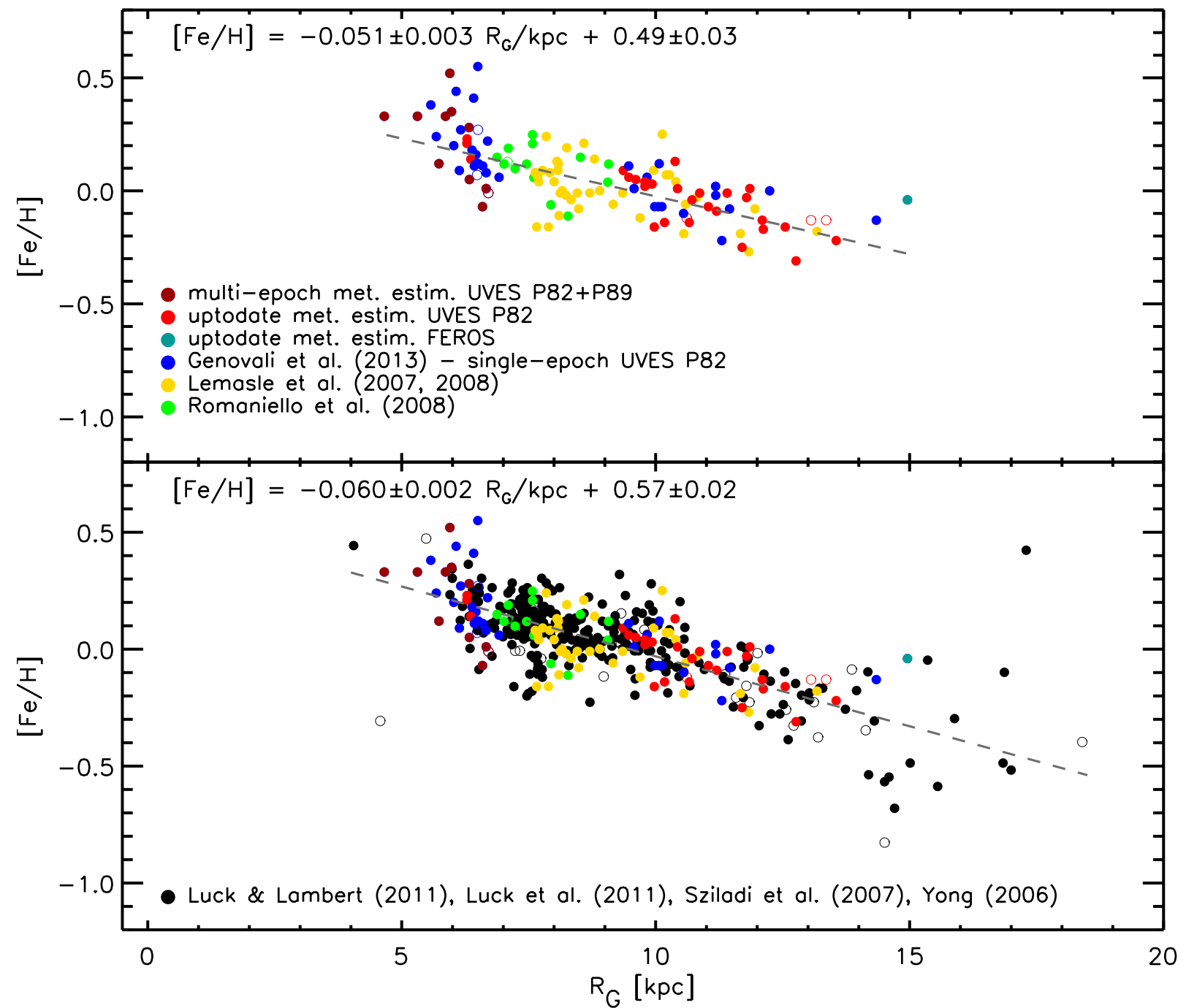
δM_K



Romaniello et al. (2008)

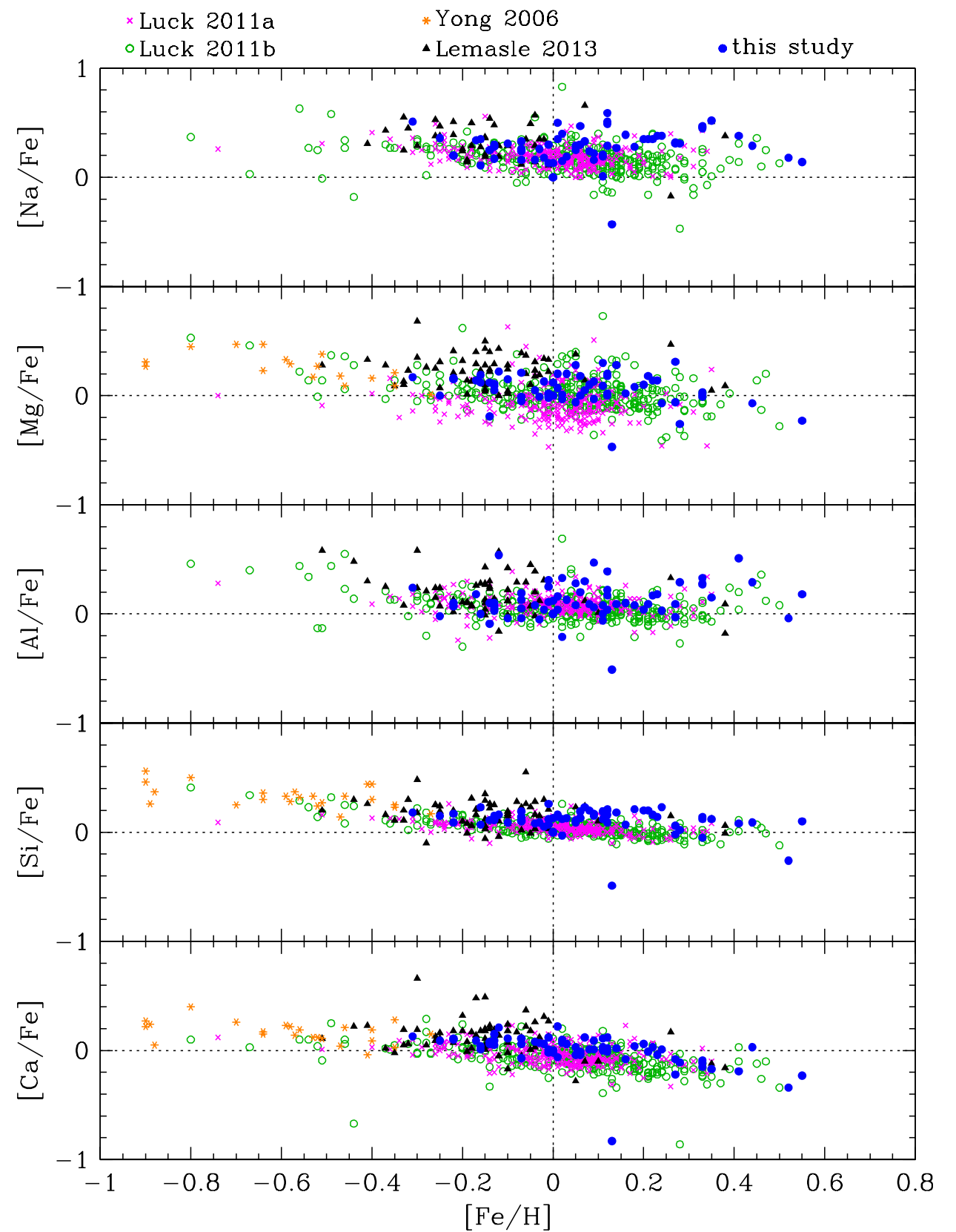
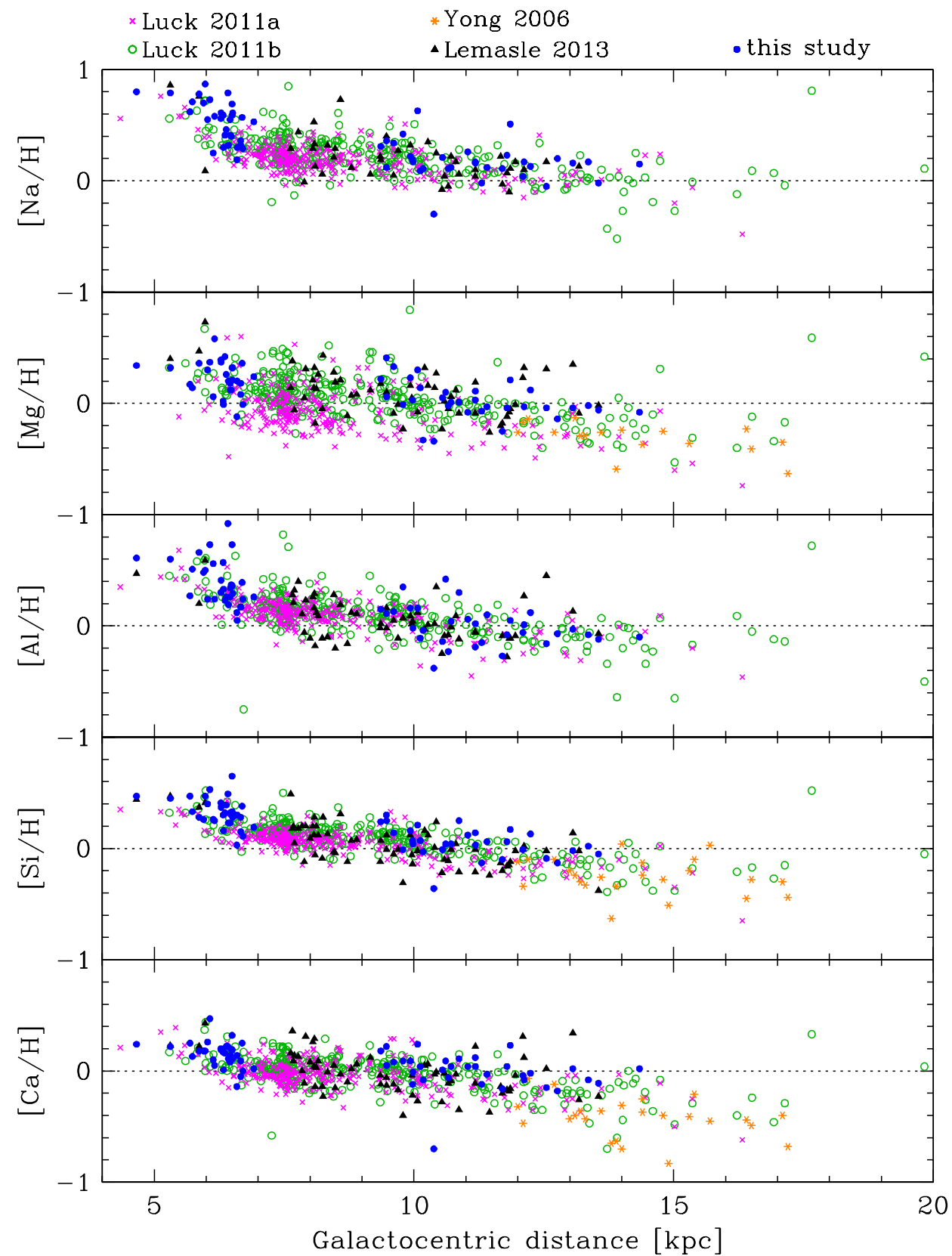
Introduction

Metallicity gradient

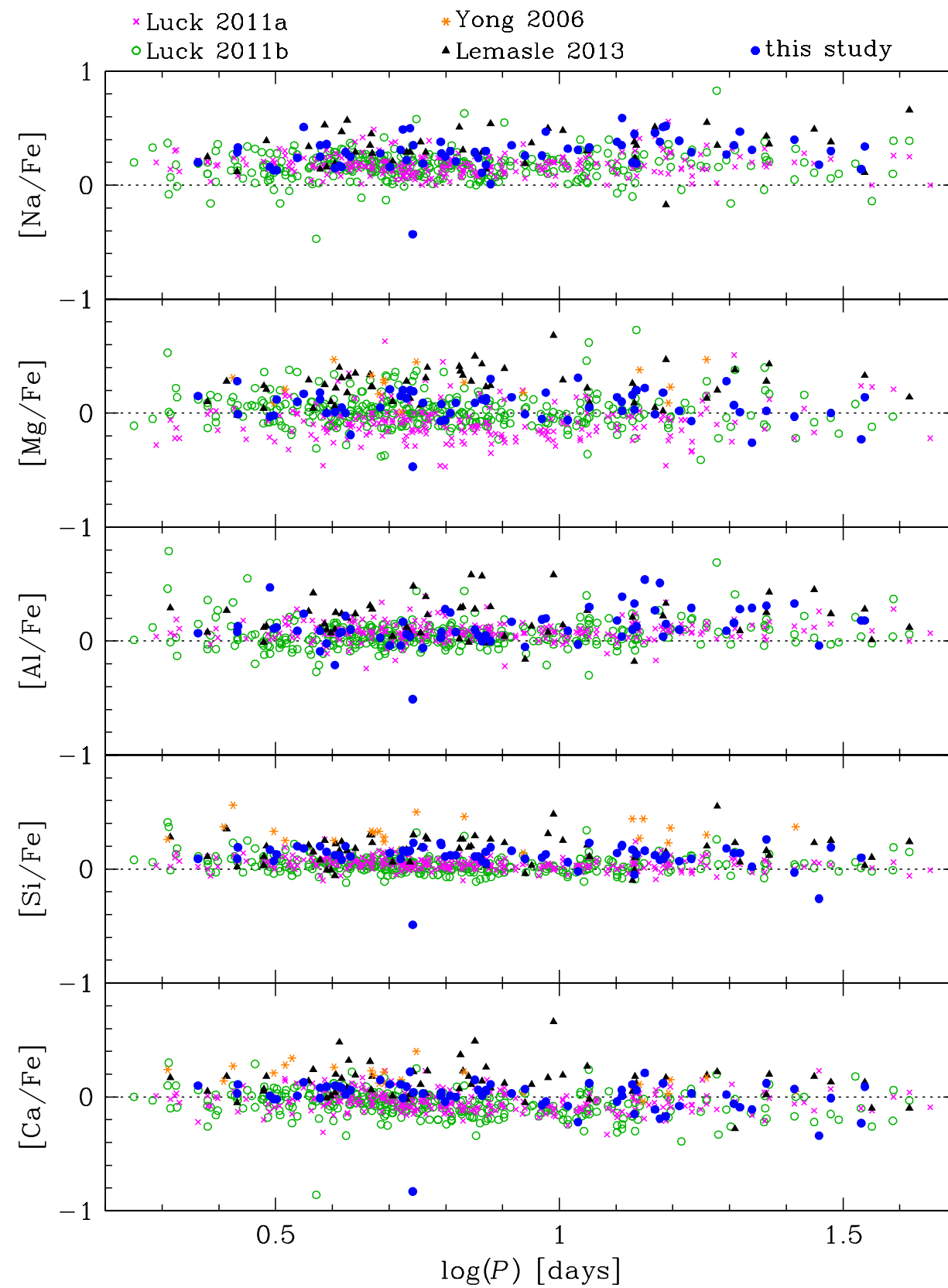


Genovali et al. (2014)

Introduction



Introduction



Introduction

- ★ Constraints on the chemical composition of extragalactic Cepheids are based on $[O/H]$ nebular abundances of HII regions

- ★ Zero point:

Strong Oxygen and $H\beta$ emission lines in extragalactic sources

X

Oxygen absorption lines in Galactic Cepheids

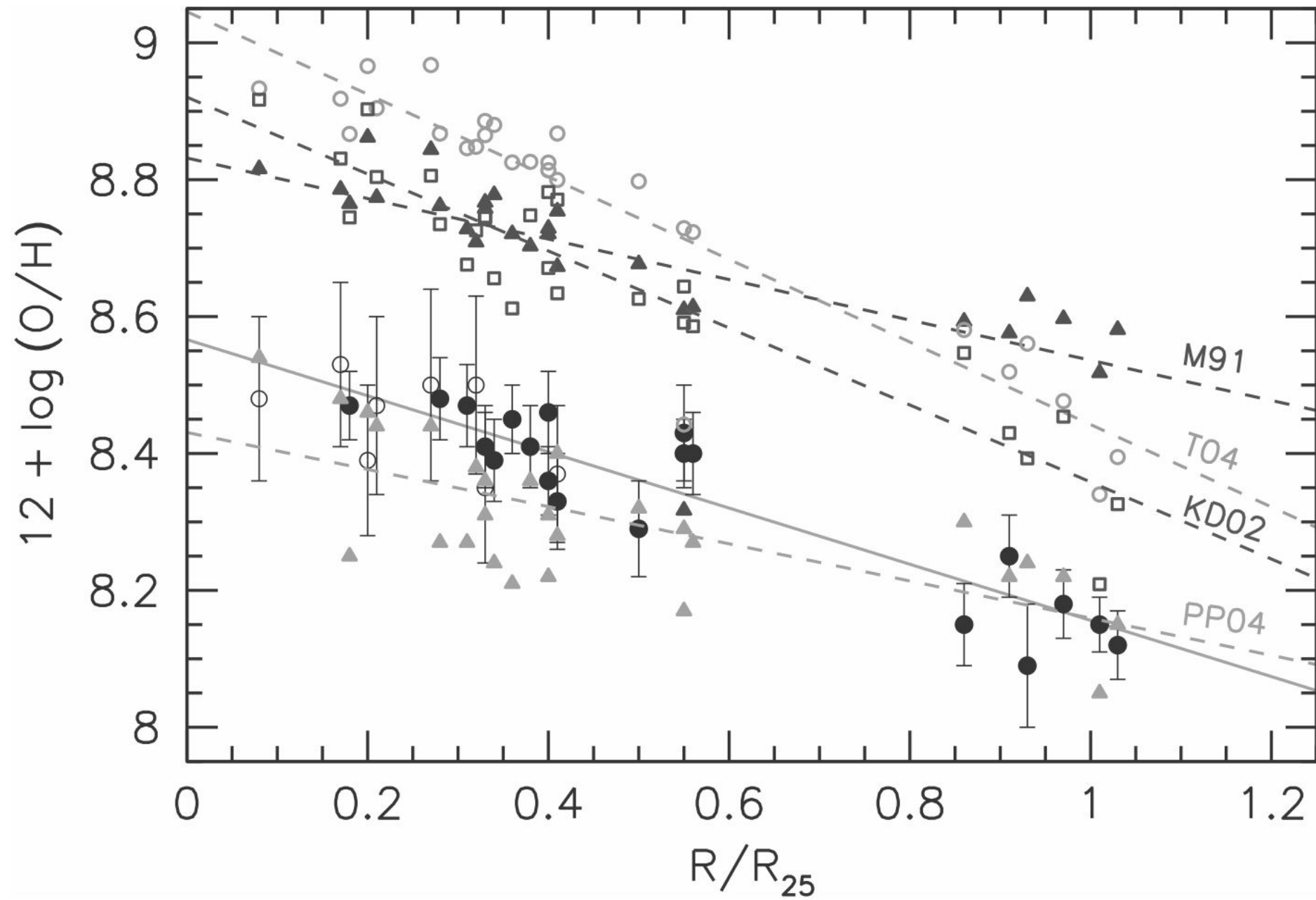
- ★ Slope:

Do $[O/H]$ in extragalactic Cepheids scale with their metal content?

- ★ This study: O, Na, Mg, Al, Si, S, Ca, and Eu in the SMC and LMC

Introduction

Oxygen abundances from H II regions in NGC 300

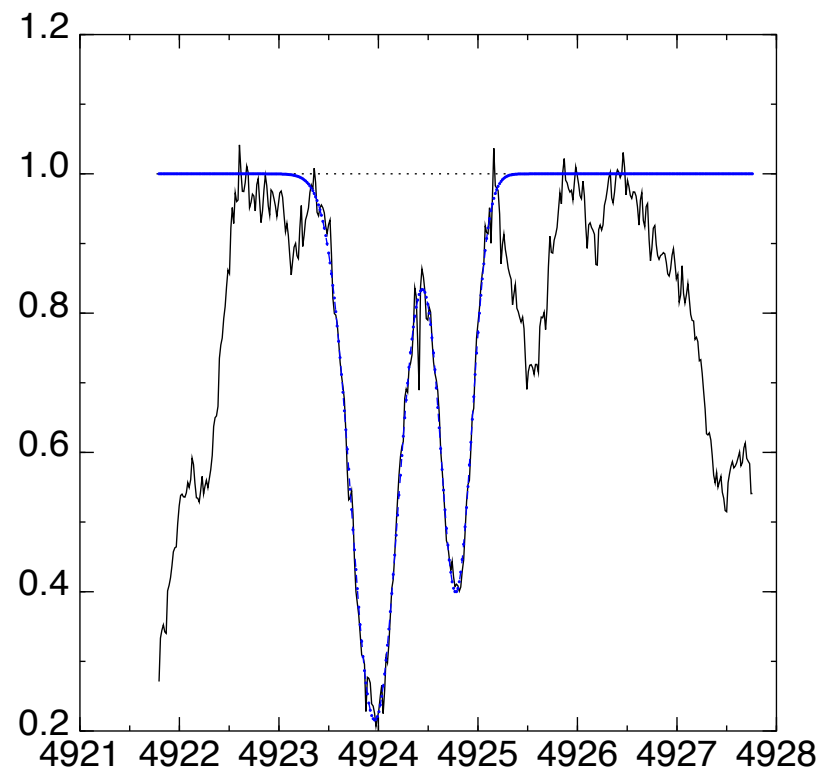


Observations and Method

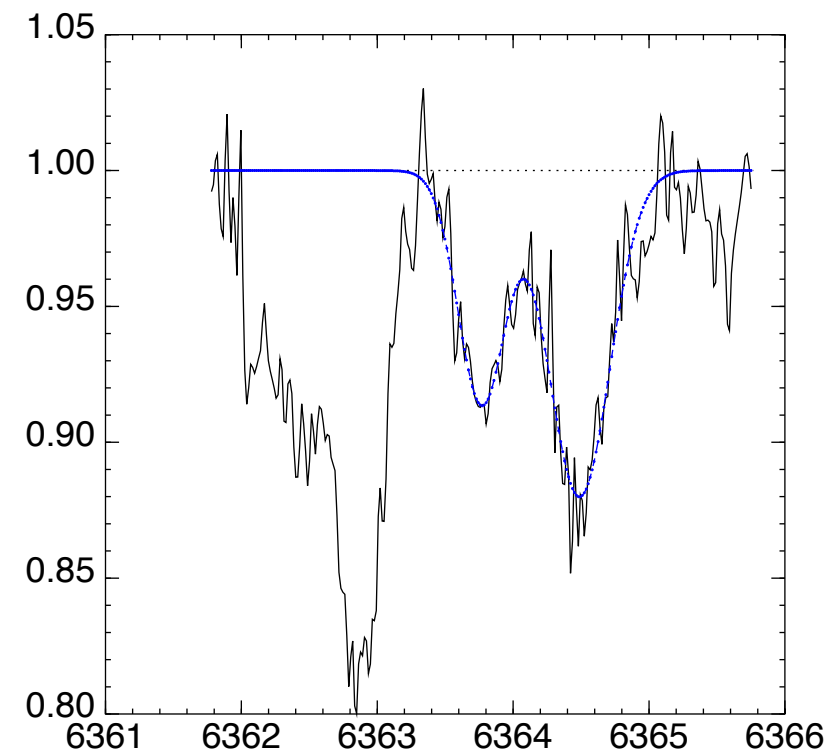
- ★ 22 stars of the Large Magellanic Cloud (LMC)
- ★ 14 stars of the Small Magellanic Cloud (SMC)
- ★ Spectra collected using UVES at VLT
 $R \sim 38000$ $S/N \gtrsim 50$ $\lambda\lambda 4800\text{-}6800 \text{ \AA}$
- ★ Photospheric parameters derived from equivalent widths (EW) of Fe I and Fe II lines
- ★ Abundances from the EWs of atomic lines

Observations and Method

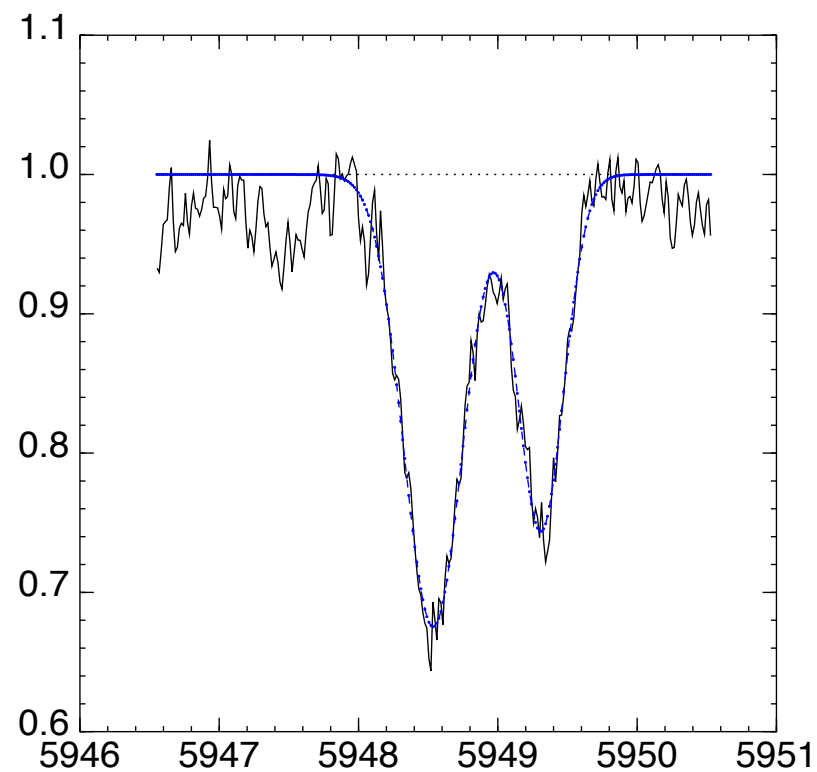
Fe I
 $\lambda 4924.77 \text{ \AA}$



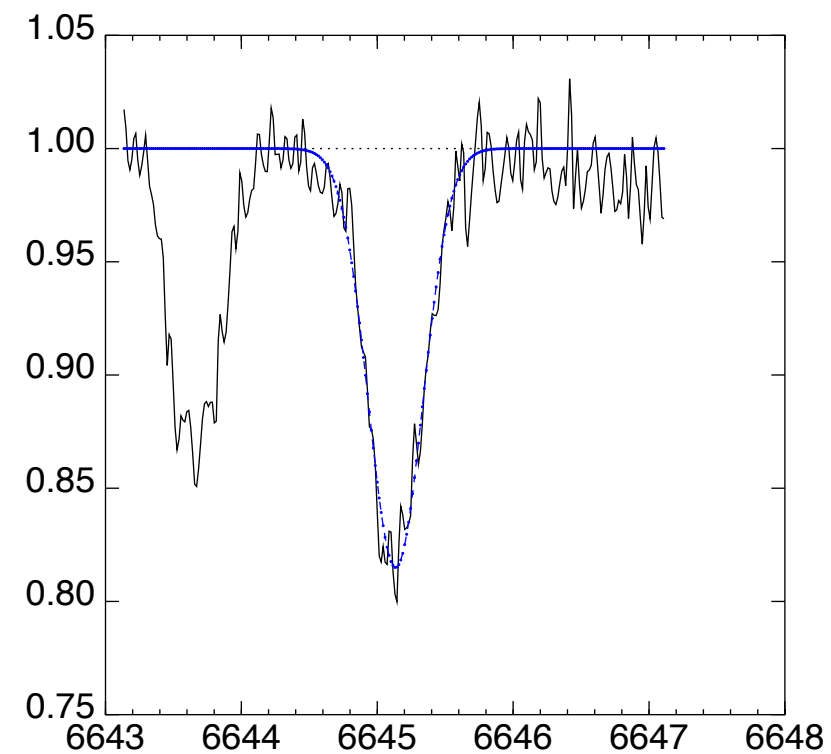
O I
 $\lambda 6363.77 \text{ \AA}$



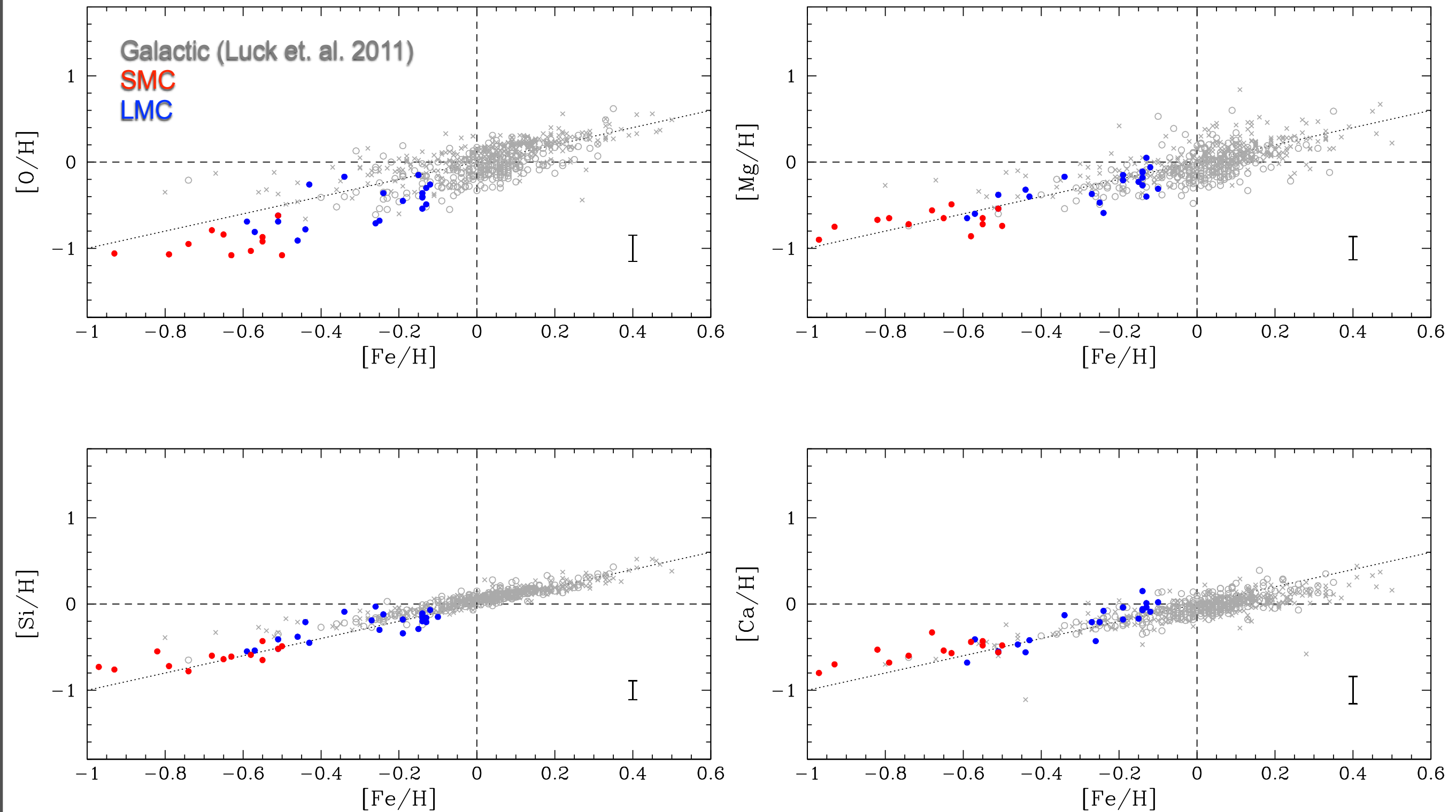
Si I
 $\lambda 5948.54 \text{ \AA}$



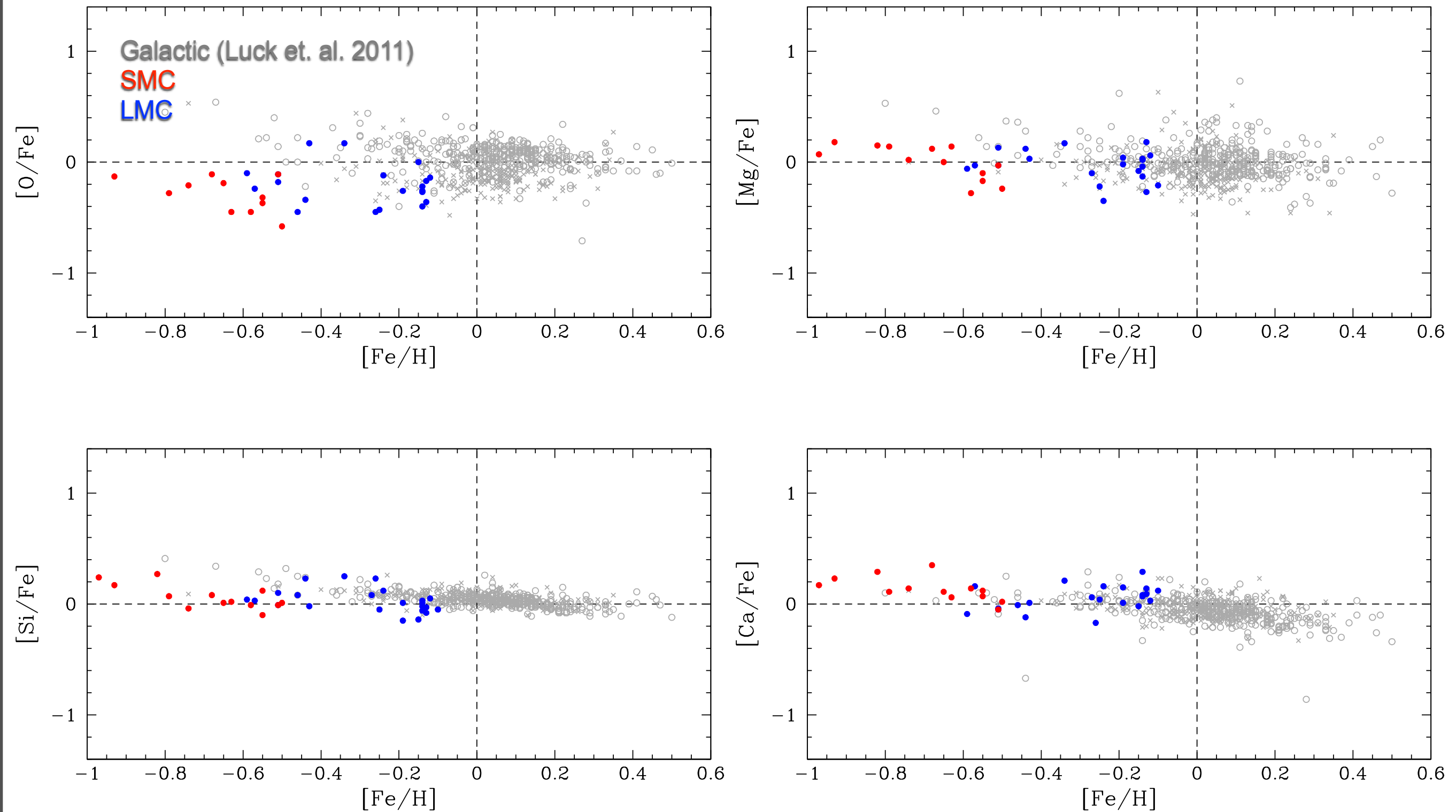
Eu II
 $\lambda 6645.13 \text{ \AA}$



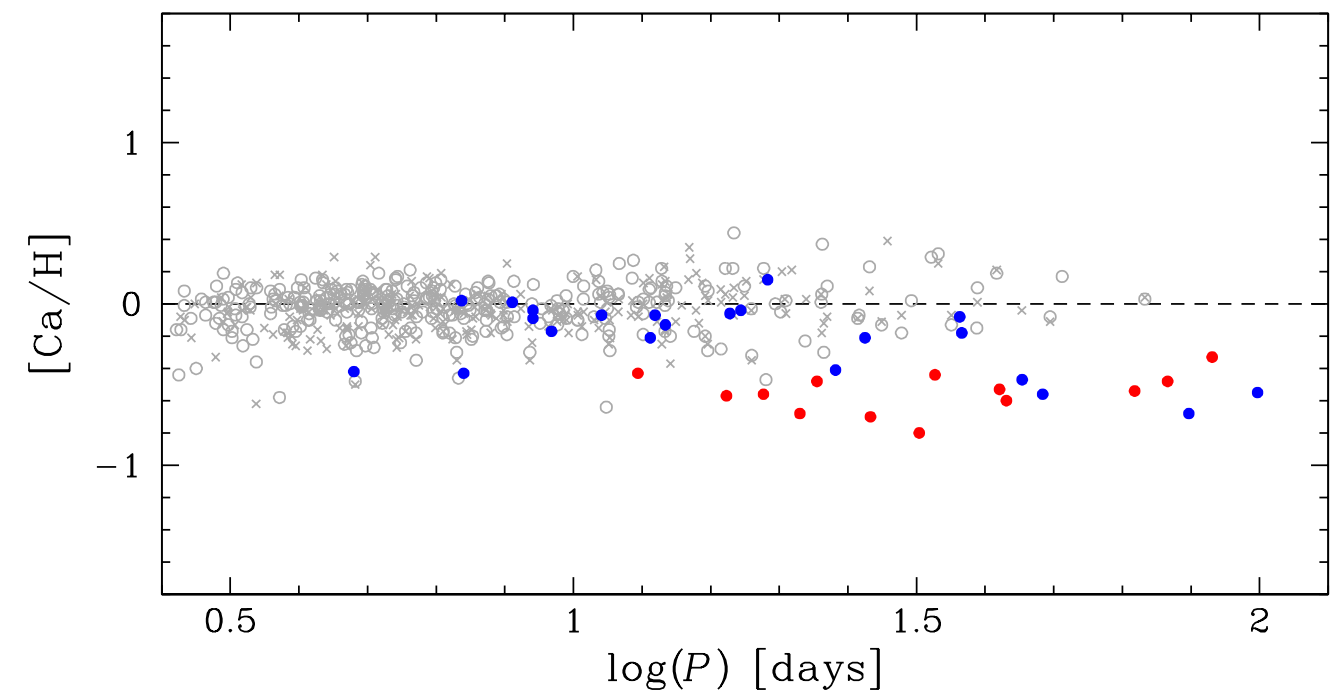
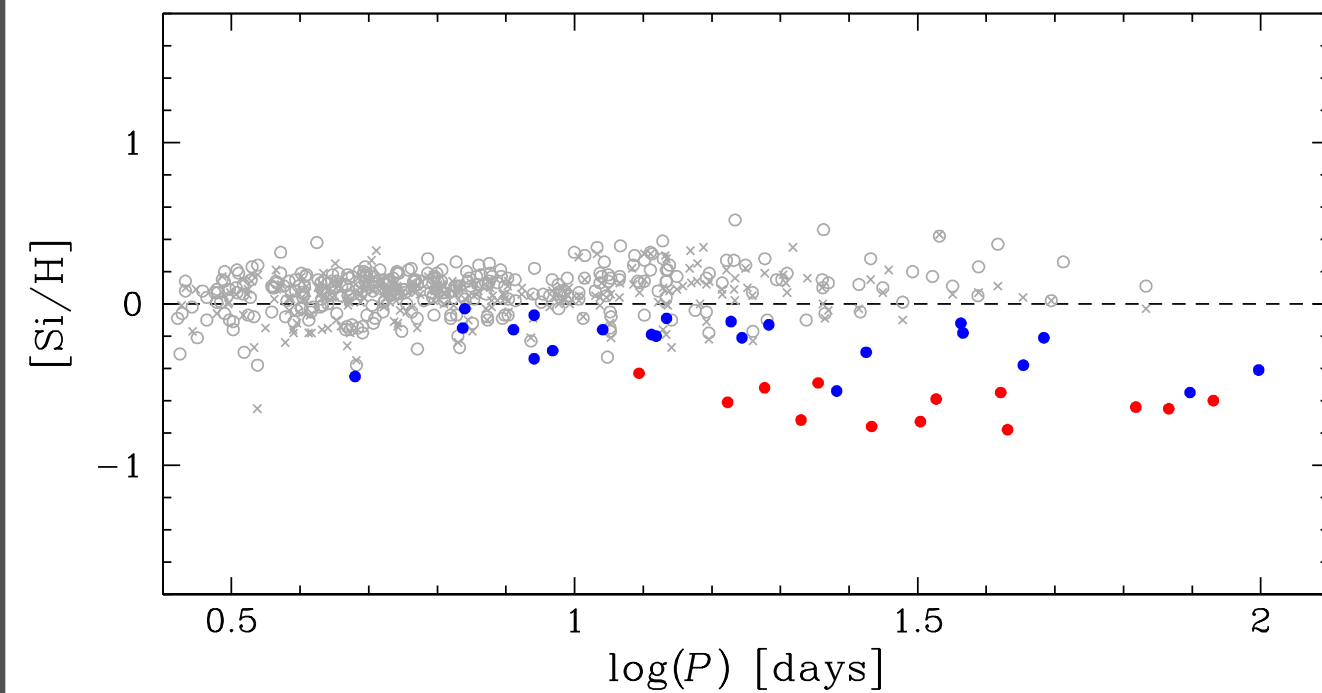
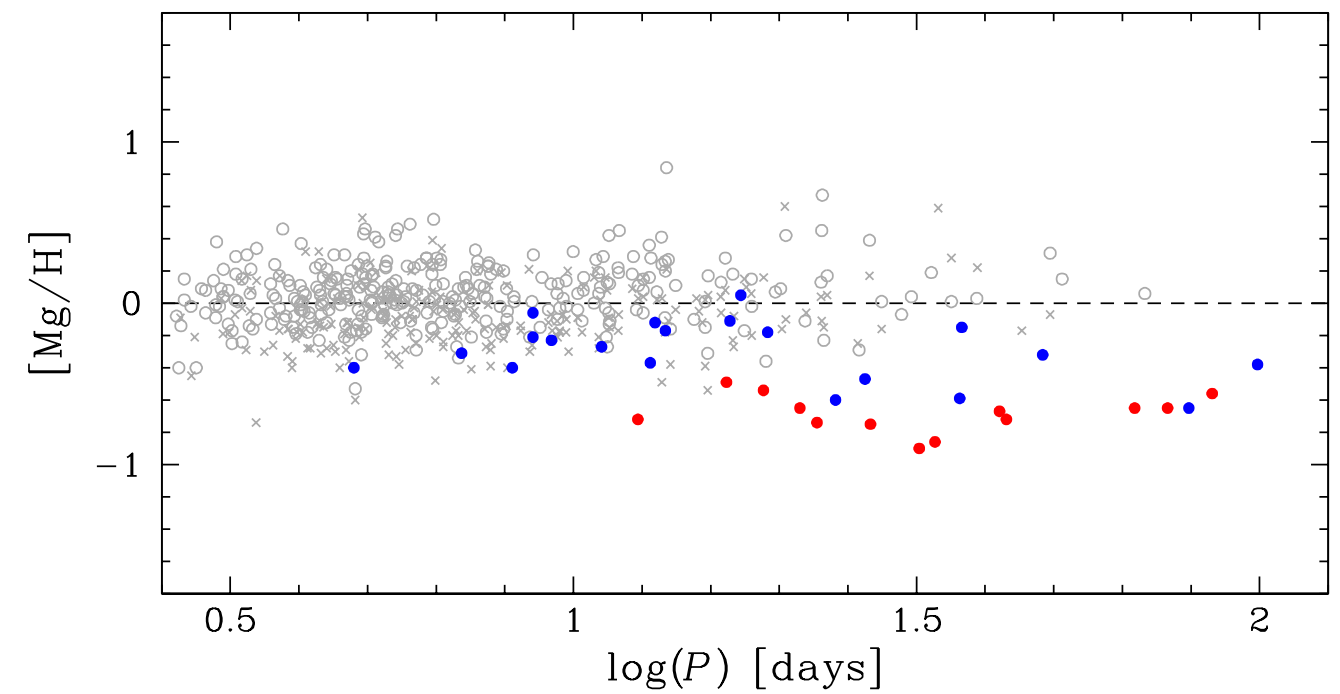
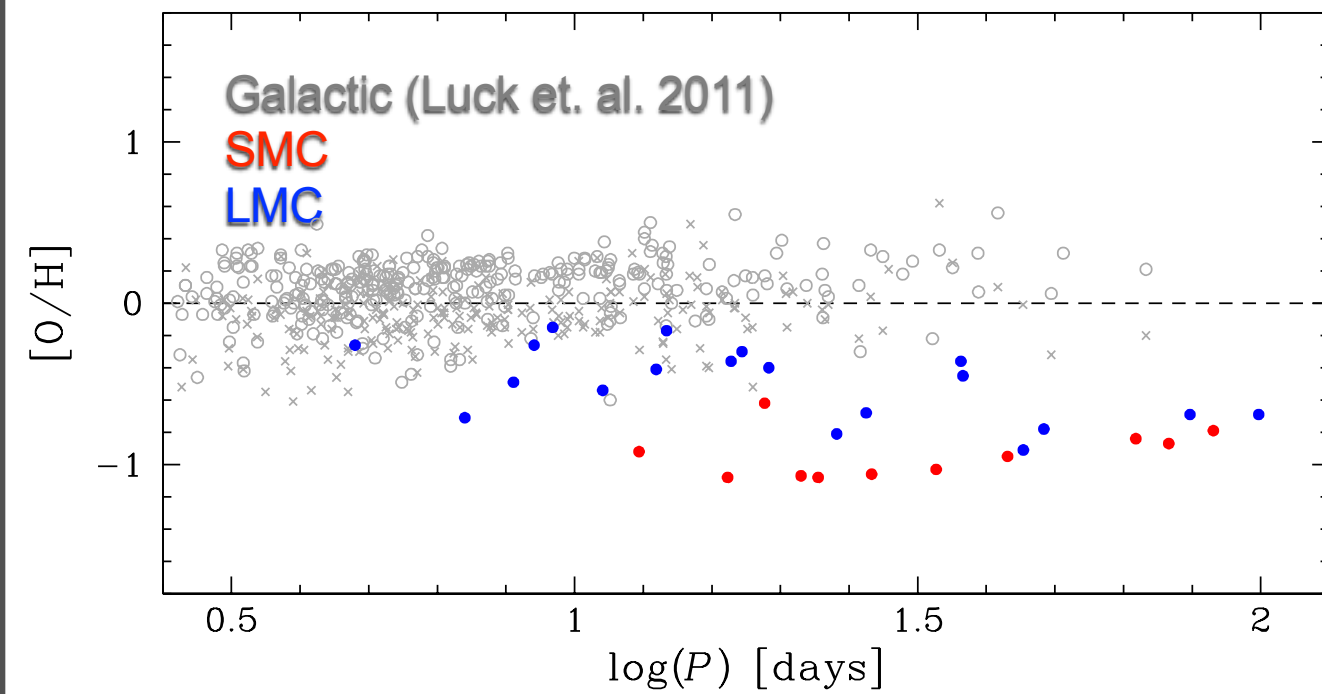
First results: $[\alpha/H]$ vs. $[Fe/H]$



First results: $[\alpha/\text{Fe}]$ vs. $[\text{Fe}/\text{H}]$



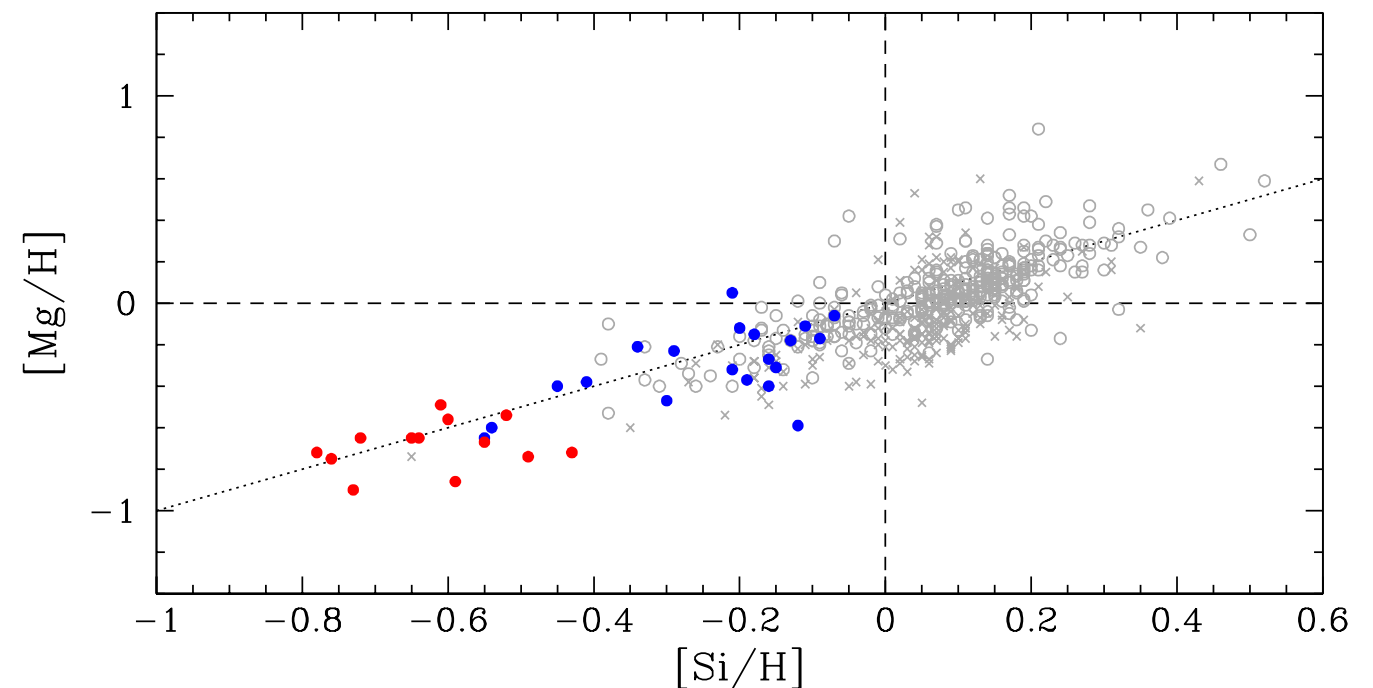
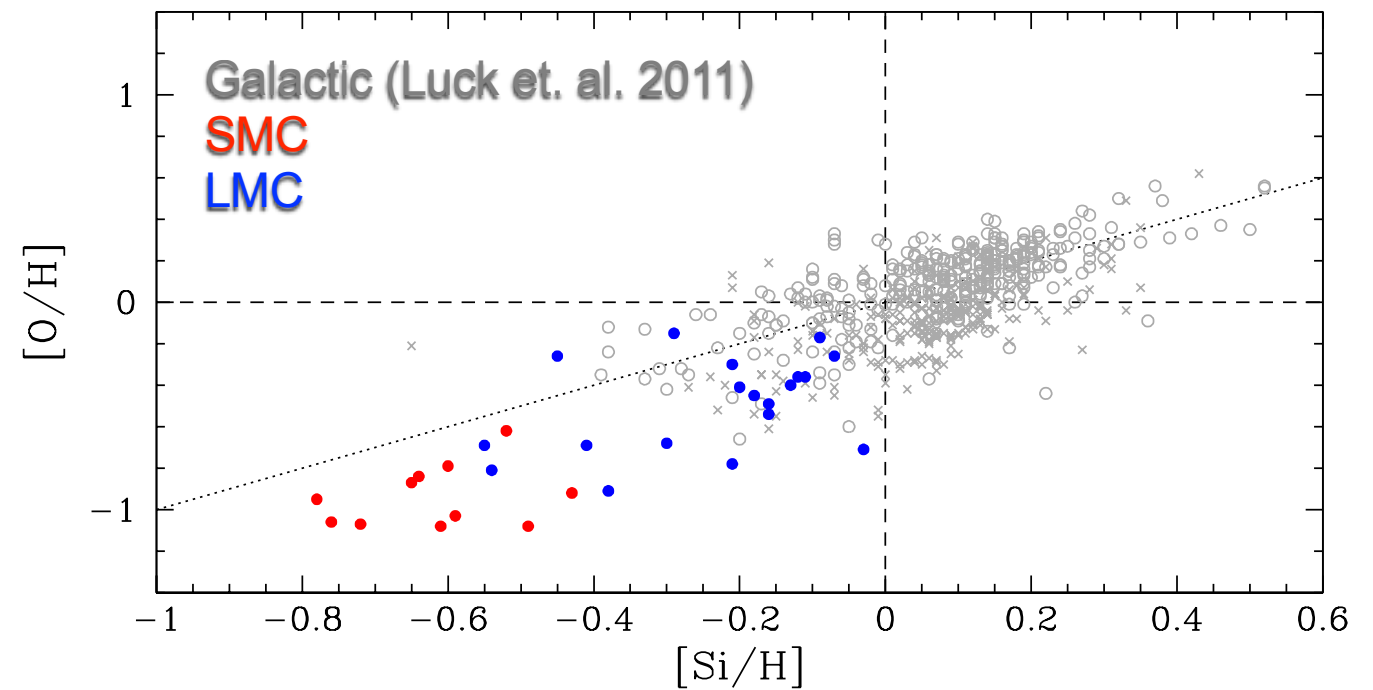
First results: $[\alpha/\text{H}]$ vs. $\log(P)$



First results: Hydrostatic vs. Explosive elements

Hydrostatic burning phases:
O, Mg, Na, Al, Cu

SN explosion event:
Si, Ca, Ti



Summary

- ★ Oxygen abundances do scale with the metal content of extragalactic Cepheids, but we found a zero-point difference
- ★ Decreasing $[\alpha/\text{H}]$ with increasing pulsation period
- ★ Hydrostatic vs. explosive elements: negative ratios in the Magellanic Clouds
- ★ E-ELT: higher magnitudes, larger distances